

not very different from Rowland's. It is doubtful if they would suit another star unless it were exactly the same type. The results from several late type stars, whose orbits have been computed at this observatory, seem to show that the apparent wave-lengths of blends, as measured on low dispersion plates, vary markedly with a slight variation in the spectral type.

The usual method of first finding approximate elements and then correcting these differentially was followed. One correction was found sufficient. The successive steps in the solution and the results are given below. In the radial velocity curve the initial date is Julian Day 2,420,500.

TABLE 1

Element	Wave-length.	Correction	Algebraic Residual.	Arithmetic Residual	Number of times Measured.
<i>Fe</i>	4005.426	- .037	- 2.8	4.4	24
<i>Fe</i>	4015.975	+ .011	+ 0.8	1.8	4
<i>Fe</i>	4063.756	+ .114	+ 8.4	8.4	4
<i>Fe</i>	4071.939	- .004	- 0.3	2.7	10
<i>Fe</i> +	4092.737	- .023	- 1.7	4.1	36
<i>V</i>	4116.643	+ .041	+ 3.0	5.9	16
<i>Fe-Ti</i>	4123.748	+ .033	+ 2.4	4.7	13
<i>Fe</i>	4127.992	- .029	- 2.1	4.1	27
<i>Fe</i> +	4132.361	+ .021	+ 1.5	3.3	22
<i>Fe</i>	4191.697	+ .011	+ 0.8	3.3	20
<i>Fe</i>	4202.291	- .032	- 2.3	5.2	33
<i>Fe</i>	4236.057	+ .013	+ 0.9	4.3	17
<i>Fe</i>	4250.616	- .013	- 0.9	3.3	29
<i>Fe</i>	4258.605	- .010	- 0.7	1.5	6
<i>Fe</i>	4271.829	+ .006	+ 0.4	3.1	40
	4275.119	- .003	- 0.2	3.3	18
<i>Fe-Ti</i>	4282.722	+ .027	+ 1.9	4.2	36
<i>Cr</i>	4289.766	+ .003	+ 0.2	3.0	34
<i>Ti</i>	4314.866	- .010	- 0.7	4.0	30
	4323.612	+ .073	+ 4.8	7.1	15
<i>Fe</i>	4325.792	- .003	- 0.2	4.8	12
<i>Fe</i>	4401.890	- .004	- 0.3	4.0	37
<i>Fe</i>	4415.225	+ .007	+ 0.5	3.9	40
<i>Fe</i>	4430.469	- .023	- 1.6	3.7	35
<i>Fe-Mn</i>	4462.037	- .029	- 2.0	4.6	27
<i>Fe-Ti</i>	4519.766	+ .015	+ 1.0	4.1	33
<i>Cr-Ti</i>	4571.891	- .009	- 0.6	4.3	39