at the beginning of 1916, some high dispersion equipment may record the fainter component spectrum. If the star should prove to be a visual binary, and we combined the spectrographic measures of the relative velocity of the two components in the line of sight with the orbit data from micrometric measures, we would have an indirect means of determining the parallax of the star. So far as is known to the writer no values of the parallax of this star by any method have been published.

The curve shown represents the elements as corrected. The plain circles represent our own observations grouped, while the circles with dark centres the grouped measures of other observatories.

It is with some hesitation that the foregoing elements are submitted. Owing to the paucity of observations around the peak of the curve we cannot regard these elements as definite, but they can be regarded, I think, as fairly approximate. It is just possible that this discussion may result in some unpublished measures of this star's radial velocity being brought to light.

A good share of the credit for this determination of the orbital elements should be given to those observatories who made the early observations, without which our own measures would be inadequate.

DOMINION OBSERVATORY,
OTTAWA, CANADA,
April, 1912.